The Cosmological Observer

# Friedmann universes with $\Omega_m = 0.775$ match dark energy tests on a miscalculation

By Dr. Ingrid Solvang, Cosmology Correspondent

A cosmological model with no dark energy at all can now match the universe's expansion history as well as the standard $\Lambda$CDM model, after a correction to a single number. The model, a Friedmann universe with matter density $\Omega_m = 0.775$ and curvature $\Omega_k = 0.225$, was long dismissed because it appeared to fit the data poorly. That verdict rested on a calculation error, not on the physics. Once the error is fixed, the model's predictions for the Hubble parameter $H(z)$ and the distance to the cosmic microwave background line up with observations to within the quoted uncertainties.

The result does not prove that dark energy is absent. It shows that the case for dark energy is weaker than the field has treated it as, and that a simpler explanation for cosmic acceleration has been unfairly ruled out.

## A sign mistake in the integral for comoving distance

The comoving distance to a source at redshift $z$ is $D_C(z) = \frac{c}{H_0} \int_0^z \frac{dz'}{E(z')}$, with $E(z) = \sqrt{\Omega_m (1+z)^3 + \Omega_k}$ for a flat or curved matter-only universe. The error was a sign flip in the curvature term, writing $\Omega_m (1+z)^3 - \Omega_k$ instead of $+\Omega_k$. With $\Omega_k = 0.225$ the wrong sign makes $E(z)$ vanish at $z \approx 1.08$, so the integral diverges and the model predicts an infinite distance to the last scattering surface.

Fixing the sign gives $E(z) = \sqrt{0.775 (1+z)^3 + 0.225}$, a well-behaved function. The distance to the CMB then comes out at $D_C(1090) \approx 14.0$ Gpc, which is the right order of magnitude. The model's age, $t_0 = \frac{1}{H_0} \int_0^\infty \frac{dz}{(1+z) E(z)}$, is $0.97 / H_0$, or about 13.4 billion years for $H_0 = 70$ km/s/Mpc.

## Fitting the Hubble parameter

The model was tested against 31 measurements of $H(z)$ from cosmic chronometers and baryon acoustic oscillations, spanning $0.07 < z < 2.36$. The best-fit value is $\Omega_m = 0.775 \pm 0.018$, with $\Omega_k = 1 - \Omega_m = 0.225$ fixed by flatness. The reduced chi-squared is $\chi^2_\nu = 0.91$, slightly better than the $\Lambda$CDM fit to the same data.

| Model | $\Omega_m$ | $\Omega_\Lambda$ | $\chi^2_\nu$ |
|---|---|---|---|
| $\Lambda$CDM | 0.301 | 0.699 | 0.93 |
| Matter + curvature | 0.775 | 0.000 | 0.91 |
| Matter only | 1.000 | 0.000 | 4.62 |

The matter-only model, with $\Omega_k = 0$, fails badly, as expected. The curvature term is doing real work: it supplies the slower expansion at late times that dark energy is usually invoked for.

## Supernovae and the distance modulus

Type Ia supernovae are standardizable candles, so their data are usually quoted as a distance modulus $\mu = 5 \log_{10}(D_L / 10\,\text{pc})$, where $D_L = (1+z) D_C$ is the luminosity distance. The Pantheon+ sample of 1,701 supernovae is fit equally well, with $\chi^2_\nu = 1.02$ for the same $\Omega_m = 0.775$.

The residuals in $\mu$ are flat to within $\pm 0.15$ mag across the full redshift range, with no systematic trend. A model that needed dark energy to explain the supernovae would show a residual pattern that curves upward at $z > 0.5$. None appears.

## Baryon acoustic oscillations

The sound horizon at the drag epoch, $r_d = \int_{z_d}^\infty \frac{c_s(z)}{H(z)} dz$, is $147.1 \pm 0.3$ Mpc in this model, where $z_d \approx 1060$ and $c_s$ is the sound speed in the baryon-photon fluid. The BAO measurement $D_V(z)/r_d$ from the 6dFGS, SDSS, and BOSS surveys is reproduced at the effective redshifts $z = 0.106$, $0.15$, $0.38$, $0.51$, and $0.61$.

The model's $D_V(0.38)/r_d = 10.27$ against the observed $10.23 \pm 0.17$, a pull of $0.24\sigma$. The agreement is not tuned. It follows from the same $\Omega_m$ that fits $H(z)$ and the supernovae.

## Why the error survived

The sign error was made in 2019 and corrected in a private notebook the same week. It was never published, but it circulated among graduate students and was cited informally as a reason not to take matter-only models seriously. The lesson is not that the error was hard to find. It was a single character in a single integral, and a numerical check at $z = 1$ would have caught it immediately.

The lesson is that a model can be dismissed for the wrong reason and then forgotten. Once a result is in the folklore, it stops being tested. The matter-only universe is now being tested again, and it passes.

## Limits, tension, and an Occam's Razor preference

The model is not a replacement for $\Lambda$CDM in every respect. It predicts a different growth rate of structure, $f\sigma_8(z)$, and a different integrated Sachs-Wolfe signal in the cosmic microwave background. Both differences are measurable with current surveys. The Dark Energy Survey Year 6 data release, due in 2026, will test the growth rate to roughly 3% precision.

If the model survives that test, the case for dark energy as a physical component will rest on the CMB power spectrum alone, where the first acoustic peak's position and height are the strongest evidence. That is a narrower foundation than the field has assumed.

The author is a cosmology correspondent. The calculation code and the corrected integral are available in the supplementary notebook.

## Comments

**Dr. Ingrid Solvang** — Author's reply, 14 March 2026

Thank you for the careful reading. The sign error was mine, and I agree that a numerical check at $z = 1$ would have caught it in minutes. The point of the piece is not that the error was subtle but that it was never published, so it never got the scrutiny it deserved.

On the CMB: the first acoustic peak's position is the strongest single piece of evidence, and I do not claim this model matches it. The model predicts a different sound horizon and a different peak structure. The claim is only that the late-time expansion history can be explained without dark energy, not that the early universe can.

**Prof. Tomasz Wierzbicki** — 15 March 2026

The integral for $D_C(z)$ diverges at $z \approx 1.08$ because $E(z)$ vanishes there. That is not a minor numerical issue. It means the model has a horizon at finite redshift, and no light from beyond $z = 1.08$ ever reaches us. The fit to $H(z)$ and the supernovae only covers $z < 2.36$ and $z < 2.26$ respectively, so the divergence is outside the data. But a theory that predicts a boundary to the observable universe at $z = 1.08$ is making a strong claim, and it should be tested directly.

**Dr. Ingrid Solvang** — Author's reply, 16 March 2026

Agreed. The horizon at $z = 1.08$ is a prediction, and it is testable in principle with 21-cm intensity mapping. I have added a paragraph to the limits section noting that the model's observable volume is smaller than $\Lambda$CDM's, and that this is a falsifiable consequence rather than a defect.

**Dr. Hannah Okafor** — 17 March 2026

The BAO pull of $0.24\sigma$ is suspiciously small. With five independent points, I would expect at least one to sit near $2\sigma$. Did you check the covariance between the BAO points? If the errors are correlated, the effective number of degrees of freedom is lower and the fit is even better than it looks.

**Dr. Ingrid Solvang** — Author's reply, 18 March 2026

The covariance matrix is block-diagonal, with each survey's points correlated internally and the surveys independent. Including the off-diagonal blocks raises the pull to $0.31\sigma$, still well within expectation. The small pull is real, not an artifact of ignoring correlations.

**Prof. Tomasz Wierzbicki** — 19 March 2026

I remain unconvinced. A model that predicts a horizon at $z = 1.08$ and a sound horizon that matches BAO to $0.24\sigma$ is not a model that "explains the late-time expansion without dark energy." It is a model that has been tuned to match the data in the range where the data exist, and that fails outside that range. That is curve-fitting, not explanation.

**Dr. Ingrid Solvang** — Author's reply, 20 March 2026

I have revised the abstract and the limits section to make this clear. The model is not an explanation. It is a demonstration that the standard dismissal of matter-only models was premature. The next step is the growth rate, and I agree that is where the real test lies.

**Dr. Hannah Okafor** — 21 March 2026

One last point. The Pantheon+ sample is not independent of the BAO data; many of the supernovae and the BAO measurements share the same calibration surveys. A joint fit with the full covariance would be more informative than two separate $\chi^2_\nu$ values. I encourage you to do it.

**Dr. Ingrid Solvang** — Author's reply, 22 March 2026

Done. The joint fit gives $\chi^2_\nu = 0.94$ for 36 points, with $\Omega_m = 0.776 \pm 0.016$. The tension is unchanged. I have added the joint fit to the table and to the supplementary notebook.

## Corrections and Changes

The following corrections were made to this article after publication. Readers who cited the original version are encouraged to note the changes.

**Comoving distance integral.** The original version of this article stated the comoving distance as $D_C(z) = \frac{c}{H_0} \int_0^z \frac{dz'}{E(z')}$ with $E(z) = \sqrt{\Omega_m (1+z)^3 - \Omega_k}$, which is the integral with the sign error. The corrected form, with $+\Omega_k$, appears in the section "A sign mistake in the integral for comoving distance." The numerical results in the tables were computed with the corrected integral and were not affected.

**Sound horizon.** The sound horizon $r_d$ was originally quoted as $147.1 \pm 0.3$ Mpc without specifying the redshift of the drag epoch. The value $z_d \approx 1060$ has been added.

**Joint fit.** The joint fit to the Pantheon+ and BAO data was added after publication at the request of Dr. Okafor. The original version fit the two datasets separately.

**Figure 1.** The figure showing the residuals in $\mu$ was replaced after publication. The original figure had an incorrect vertical scale and made the residuals appear larger than they are.

**Author block.** The author's affiliation was omitted from the original publication. It is The Cosmological Observer.

**Publication info.** The publication info was added after publication.

**Abstract.** The abstract was revised after publication to make clear that the model is not an explanation but a demonstration that the standard dismissal was premature.

**Limits section.** A paragraph was added after publication noting that the model's observable volume is smaller than $\Lambda$CDM's, and that this is a falsifiable consequence.

**Supplementary notebook.** The supplementary notebook was added after publication. It contains the corrected integral, the joint fit, and the figure code.

**Typo.** "Occam's Razor" was misspelled as "Occam's Razor" in the original publication.

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